Planning an EXES observation

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2. Example: Planning a High-Resolution observation of a point source
   1. Atmospheric transmission and Doppler shifts
   2. Using the EXES ETC
EXES, the ECHELON X (cross) Echelle Spectrograph

- Echelon (high resolution) + Echelle (cross dispersion+low resolution)
- R=1,000-100,000, i.e., 3 km/s-300 km/s
- 4.5-28.3 um.
- 4 observing modes: HIGH-MEDIUM, HIGH-LOW, MEDIUM and LOW
Medium and Low modes:

**Medium resolution mode**

R=8000-25000

~0.7 % coverage

Point source sensitivity:
5-15 Jy for S/N ~50 in 1hr

**Low mode example**

*SHARED RISK in Cycle 9*, due to high background and fringing

R=1000- 5000

~4% coverage

Point source sensitivity:
2-10 Jy for S/N ~50 in 1hr
High-medium and High-Low modes

(a) HIGH-MEDIUM
In general: $R = 50,000-100,000$

Point source sensitivity,
$20-40$ Jy for $S/N \sim 50$ in $1$hr

Example above: $13.7 \text{um}$, $4^{\text{th}}$ order
$20''$ slit length
$R=5.4 \text{ cm}^{-1}$ coverage: $0.7%$

(b) HIGH-LOW
In general, $R = 50,000-100,000$

Point source sensitivity
$30-80$ Jy for $S/N \sim 50$ in $1$hr

Example (b):
$13.8 \text{um}$
$2''$ slit length (PSF $\sim 2.7''$)
$41 \text{ cm}^{-1}$ coverage: $5.8%$
Example Observation:
Resolve water vapor abs. in a protostar

Example observation:
ISO/SWS detected abundant gas-phase water toward the massive YSO, AFGL 2591

But $R \sim 2000$

Insufficient for measuring line widths or resolving complex profiles

Resolved line profiles will reveal the location and perhaps the chemical origin of the H$_2$O

![Graph of GL 2591 H$_2$O $\nu_2$](chart)

**Fig. 2.** From top to bottom: Normalized ISO–SWS spectrum observed toward GL 2591; theoretical spectra with $T_\text{ex} = 300$ K, $b_\text{D} = 7.5$ km s$^{-1}$; $T_\text{ex} = 38$ K, $b_\text{D} = 5.3$ km s$^{-1}$; and $T_\text{ex} = 1000$ K, $b_\text{D} = 7.8$ km s$^{-1}$, shifted by $-0.2$, $-0.4$, and $-0.6$ respectively. All model spectra have $N$(H$_2$O) $= 2 \times 10^{18}$ cm$^{-2}$.

(Helmich et al., 1996)
Example observation:
resolve H2O toward AFGL 2591

Specific goals:

- obtain S/N~100, R=86,000 spectrum of the massive protostar AFGL 2591 covering 6.07 – 6.12 um. About a strong dozen ro-vib H2O lines.

- Detect ν2 1_1,1-0_0,0 ground state transition of para-H2O at 1634.9670 cm⁻¹ (6.1163um) for sensitivity to cold gas
2. Preparing EXES Observation: ATRAN & Doppler Shifts

Use ATRAN: [https://atran.sofia.usra.edu/cgi-bin/atran/atran.cgi](https://atran.sofia.usra.edu/cgi-bin/atran/atran.cgi)

ATRAN: enter typical values
altitude: 41000 ft
Zenith angle: 45 deg

At rest velocity, this is hopeless...

Need a Doppler shift!
2. Preparing EXES Observation: ATRAN & Doppler Shifts

Use ATRAN: [https://atran.sofia.usra.edu/cgi-bin/atran/atran.cgi](https://atran.sofia.usra.edu/cgi-bin/atran/atran.cgi)

Telluric line has HWHM~31 km/s

![Graph showing atmospheric transmission with peaks at 6.11 and 6.16 microns, labeled 62 km/s.](image)
2. Preparing EXES Observation: Doppler Shifts

Velocity of line absorption on a given date, $V_{DOP}$, taking into account velocity AFGL 2591 ($V_{LSR}$ or $V_{HELIO}$) as well as $V_{EARTH}$ in LSR or HELIO reference frame toward position AFGL 2591. Earth orbital around sun at ~30 km/s.

See details in next slide:

AFGL 2591: $V_{LSR} = -5.5$ km/s (submm CO lines) $\rightarrow V_{HEL} = -23.5$Km/s

$V_{DOP} = -34.7$ km/s on April 1 (=0.0007 μm) $\rightarrow$ $-12.6$ km/s on Oct 1

Best chance when maximally blueshifted!

Derive acceptable Doppler shift and set time constraints on observation in proposal. Tight constraints limit chances for observation to be scheduled!

Note: if line entirely free of telluric absorption, it may be better done from ground!
2. Preparing EXES Observation: Doppler Shifts

In IDL:

- Example given in baryvel.pro in IDL astronomy library:
  ```idl```
jdcnv, year, month, day, hour, jd ; convert date to Julian date
baryvel, jd, epoch, vh, vb ; heliocentric velocity of earth for given date in km/s
```idl```

- Project earth velocity toward star. RA and Dec stellar position in radians
  \[ V_{\text{EARTH}} = vh[0] \cdot \cos(\text{Dec}) \cdot \cos(\text{RA}) + vh[1] \cdot \cos(\text{Dec}) \cdot \sin(\text{RA}) + vh[2] \cdot \sin(\text{Dec}) \]

- Add radial heliocentric velocity of star to radial heliocentric velocity of the earth at that date. The sign of \( V_{\text{EARTH}} \) is negative!
  \[ V_{\text{DOP}} = V_{\text{HELIO}} - V_{\text{EARTH}} \]

- Note: to convert \( V_{\text{LSR}} \) to \( V_{\text{HEL}} \) use helio2lsr.pro
  [https://people.ok.ubc.ca/erosolo/idl/lib/helio2lsr.pro](https://people.ok.ubc.ca/erosolo/idl/lib/helio2lsr.pro)

- AFGL 2591 on April 4, 2014. \( V_{\text{dop}} = -34 \text{ km/s} \)
2. EXES Exposure Time Calculator

The EXES “Exposure Time Calculator” (ETC)
https://dcs.arc.nasa.gov/proposalDevelopment/SITE/index.jsp
http://irastro.physics.ucdavis.edu/exes/etc/etc.html

Welcome to the SOFIA - EXES Exposure Time Calculator

VERY IMPORTANT: For SOFIA proposals, enter the CLOCK time, not the integration time, as overheads depend on observing details and are captured in the ETC clock time.

Step 1

Enter either the rest-frame wavelength OR the rest-frame wavenumber to be observed: 1634.96 [4.5 - 28.5 micron, or 350 - 2220 cm⁻¹]

Check here to apply a Doppler shift: ☑ and enter the velocity: -34.7 [km/s, negative if the source is approaching]

Note that observations of features near strong telluric features can change dramatically with Earth’s orbital motion. One available tool is (link should open new window or tab) GBT’s VLSR Calculator. To use, add the source VLSR to the correction calculated for a given date.

Step 2

Next, select the instrument mode from the options below:
- Cross-dispersed High-Medium
- Cross-dispersed High-Low
- Single-order Long Slit Medium
- Single-order Long Slit Low

Click the submit button to continue on to the next step: Submit

Click here for the ETC user manual & documentation.
2. Preparing EXES Observation: Exposure Time Calculator

Slit width sets the resolution. Narrower slits block more star light (SOFIA PSF ~2.7”). Trade off between resolving power and S/N!

### Step 4 - Select a slit width

<table>
<thead>
<tr>
<th>Slit Width (arcsec)</th>
<th>Ext. Source Aperture (Slit Width x IQ, arcsec^2)</th>
<th>R</th>
<th>R</th>
<th>R</th>
<th>R</th>
</tr>
</thead>
<tbody>
<tr>
<td>1.44</td>
<td>4.77</td>
<td>112000</td>
<td>112000</td>
<td>112000</td>
<td>112000</td>
</tr>
<tr>
<td>1.89</td>
<td>6.24</td>
<td>85590</td>
<td>85590</td>
<td>85590</td>
<td>85590</td>
</tr>
<tr>
<td>2.43</td>
<td>8.01</td>
<td>66667</td>
<td>66667</td>
<td>66667</td>
<td>66667</td>
</tr>
<tr>
<td>3.23</td>
<td>10.68</td>
<td>50000</td>
<td>50000</td>
<td>50000</td>
<td>50000</td>
</tr>
</tbody>
</table>
2. Preparing EXES Observation: Exposure Time Calculator

Cross disperser grating order sets the echelon order separation, and thus the number of echelon orders (i.e., wavelength coverage) that fit on the array. Slit length is matched to the echelon order separation:

### Step 3 - Select an observing order

<table>
<thead>
<tr>
<th>Order</th>
<th>Grating Angle (alpha)</th>
<th>R (with default slit)</th>
<th>Minimum Wavelength</th>
<th>Maximum Wavelength</th>
<th>Minimum Wavenumber</th>
<th>Maximum Wavenumber</th>
<th>Slit Length</th>
<th>Point Source Nodding</th>
</tr>
</thead>
<tbody>
<tr>
<td>6</td>
<td>32.854</td>
<td>112000</td>
<td>6.06134</td>
<td>6.17088</td>
<td>1620.51</td>
<td>1649.8</td>
<td>3.75</td>
<td>Must be off-slit.</td>
</tr>
<tr>
<td>7</td>
<td>39.63</td>
<td>112000</td>
<td>6.07295</td>
<td>6.15889</td>
<td>1623.67</td>
<td>1646.65</td>
<td>5.06</td>
<td>Must be off-slit.</td>
</tr>
<tr>
<td>8</td>
<td>47.192</td>
<td>112000</td>
<td>6.08283</td>
<td>6.14877</td>
<td>1626.34</td>
<td>1643.97</td>
<td>6.9</td>
<td>Must be off-slit.</td>
</tr>
<tr>
<td>9</td>
<td>56.118</td>
<td>112000</td>
<td>6.09202</td>
<td>6.1394</td>
<td>1628.82</td>
<td>1641.49</td>
<td>10.01</td>
<td>Must be off-slit.</td>
</tr>
</tbody>
</table>
2. Preparing EXES Observation: Exposure Time Calculator

Cross disperser grating order sets the echelon order separation, and thus the number of echelon orders (i.e., wavelength coverage) that fit on the array. Slit length is matched to the echelon order separation and thus whether on-slit nodding is possible:

### Step 3 - Select an observing order

<table>
<thead>
<tr>
<th>Order</th>
<th>Grating Angle (alpha)</th>
<th>R</th>
<th>Minimum Wavelength</th>
<th>Maximum Wavelength</th>
<th>Minimum Wavenumber</th>
<th>Maximum Wavenumber</th>
<th>Slit Length</th>
<th>Point Source Nodding</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td>(Degrees)</td>
<td>(micron)</td>
<td>(micron)</td>
<td>(cm^-1)</td>
<td>(cm^-1)</td>
<td>(arcsec)</td>
<td></td>
<td></td>
</tr>
<tr>
<td>6</td>
<td>32.854</td>
<td>112000</td>
<td>6.06134</td>
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<td>56.118</td>
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<td>1628.82</td>
<td>1641.49</td>
<td>10.01</td>
<td>Must be off-slit.</td>
</tr>
</tbody>
</table>
2. Preparing EXES Observation: Clock Time and S/N

Step 6 - Enter the desired S/N per Nyquist sampled resolution element and the source properties

Enter the desired signal to noise ratio: 100

Note: The S/N ratio entered here is the target S/N within a Nyquist-sampled resolution element centered on the target wavelength. As the slit is oversampled by the detector, this assumes binning to 2 effective pixels per slit width.

<table>
<thead>
<tr>
<th>Source Type</th>
<th>Source flux/surface brightness</th>
</tr>
</thead>
<tbody>
<tr>
<td>Point Source</td>
<td>400 [Jy]</td>
</tr>
<tr>
<td>Extended Object</td>
<td>[Jy/arcsec^2]</td>
</tr>
</tbody>
</table>

Note: The source flux or surface brightness entered should be that quantity at the target wavelength.
# 2. Preparing EXES Observation: Clock Time and S/N

## Observation Summary

<table>
<thead>
<tr>
<th>Signal to noise ratio:</th>
<th>100</th>
</tr>
</thead>
<tbody>
<tr>
<td>(binned to 2 pixels per spectral resolution element)</td>
<td></td>
</tr>
<tr>
<td>Source type:</td>
<td>Point source</td>
</tr>
<tr>
<td>Source flux:</td>
<td>400 Jy</td>
</tr>
<tr>
<td>Atmosphere:</td>
<td>41,000 ft altitude, 45 degrees elevation angle</td>
</tr>
</tbody>
</table>

## Exposure Time Calculation

<table>
<thead>
<tr>
<th>EXES Clock time:</th>
<th>754.89 seconds</th>
</tr>
</thead>
<tbody>
<tr>
<td>Integration efficiency:</td>
<td>0.24</td>
</tr>
<tr>
<td>Source count rate (e-/s):</td>
<td>1421</td>
</tr>
<tr>
<td>Background count rate (e-/s):</td>
<td>16930</td>
</tr>
</tbody>
</table>

In USPOT, always use "clock time", which includes observing overheads, and including the time spent for off-source sky integration.

USPOT will add the calibration overheads (3 min) and appropriate time to acquire and peak up on the target.
(totol +5-18 min)
2. Preparing EXES Observation: Clock Time and S/N

At expected line position, S/N=100; but much better elsewhere.
For AORs in USPOT, enter the clock time
The real data

R=86,000 High_Medium EXES spectrum massive YSO AFGL 2591
Indriolo et al. 2015ApJ...802L..14I
Contact the instrument team for advice...

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